A One Parameter Family of
Expanding Wave Solutions
of the
Einstein Equations
that induces an
Anomalous Acceleration
into the
Standard Model of Cosmology
Blake Temple, UC-Davis
Collaborator: J. Smoller (PNAS: August 2009)

### Expanding wave solutions of the Einstein equations that induce an anomalous acceleration into the Standard Model of Cosmology

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According to the new proposition, the universe is not accelerating, as observations suggest. Instead, an expanding wave flowing through space-time has caused distant galaxies to appear to be accelerating away from us. This big wave, initiated after the Big Bang that is through to have sparked the universe, could explain they objects today accease to be farther away from us than they should be according to the Standard	ADVERTISEMENT
Model of cosmology. "Write saying that maybe the resulting expanding wave is actually causing the anomalous acceleration,"	Cosmetic World on Its Head       Cosmetic World on Its Head       New Food Sprinkle Convinces the
said back temple of the University of California. Laws. "We're saying that dark energy may not really be the correct explanation."	Brain to Stop Over-Eating Who Gets to Use Unsold Cruise
The researchers derived a set of equations describing expering waves that the listen's theory of general relativity, and which could also account for the apparent acceleration. Temple outlines the new idea with Joel Smoller of the University of Michigan in the Aug. 17 issue of the journal Proceedings of the National Academy of Sciences.	Cabins at Huge Discounts
While more research will be needed to see if the idea holds up, "the research could change the way astronomers view the composition of our universe," according to a summary from the journal.	sponsored links Refinance Now at 4.37% Fixed \$160,000 Mortgage for \$633mo. Free. No oblgation. Get 4 quotes.
To convince other cosmologists, the new model will have to pass muster with further inquiry.	MortgageRefinance.LendGo.com
There are many observational tests of the standard cosmological model that the proposed model must pass, aside from the tale phase of accelerated expansion," said Avi Loeb, director of the Institute for Theory and Computation at the Harvard Smithsorian Center for Astrophysics. "These include this phang	Car insurance cuotes Online Compare auto insurance quotes from top companies online. www.insurance.com
nucleosynthesis, the quantitative details of the microwave background anisotropies, the Lyman-abpta forest, and galaxy surveys. The authors do not discuss how their model compares to these tests, and whether the number of fike parameters they requere in order to fit these observational constants is smaller than in the standard model. Until they do so, it is not clear why this alternative model should be regarded as advantageous:	No Exam Life Insurance Get Super Charactan Colline – No Medical Exam Required. No-Exam-Life-Insurance org
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Johns Hopkins University astrophysicist <u>Mario Livio</u> agreed that to be seriously considered, the model must be able to predict properties of the universe that astronomers can measure.	Swine flu concerns
He said the real test 1s in whether they are able to reproduce all the observed cosmological parameters (as determined, e.g. by a combination of the Hubble Constant and the parameters determined by the CMB	www.s an aquant and we it be in the vaccine? * More from ABC News







egghead			
A Big Wave after the Big Bang? August 18th, 2009 @ 2:06 pm by andy			
Mathematicians <u>Blake Temple</u> from UC Davis an <u>published a new theory to explain why the univ</u> without invoking "dark energy."	d Joel Smoller from the University of Michigan <u>have</u> erse appears to be expanding at an accelerating pace,		
About a decade ago, astronomers realized that appears to be speeding up. To explain this, they that pushes the galaxies apart. No one knows w of energy that bubbles out of the fabric of space that it should make up about 70 percent of the dark matter, which is nearly as mysterious: mail trivial portion of the universe).	the universe is not only expanding — the expansion y came up with the concept of dark energy: a force what dark energy actually is; one idea is that is a sort a si te expands. Physicists' calculations, though, show universe. (Roughly another 30 percent is made of the energy that we can feel and touch make up a		
Temple and Smoller though, have a different ex they ought to be. A "big wave," started after th spreading out through space, pushing the galax	planation for why the galaxies are further apart than e Big Bang at the beginning of the universe, is ies apart.		
"We're saying that maybe the resulting expandi acceleration," <u>Temple told Space.com</u> .	ng wave is actually causing the anomalous		
Several other cosmologists quoted by Space.cor explain all the aspects of the known universe, a astronomers and physicists.	n were sceptical, noting that the new theory needs to nd make predictions that can be checked by		
The paper is published in the Aug. 17 issue of the	ne Proceedings of the National Academy of Sciences.		
For an overview of physics, cosmology and dark	energy, read this.		



### Quote:

Experts in advanced mathematics have recently proposed a new model to explain our Universe that is so different from what we have held as true thus far, that it has left many gasping for air. According to the new theory, it may be that our Universe is not expanding at all. Rather, galaxies appear to be pushing away from each other on account of a Big Bang-triggered phenomenon aptly named the Big Wave, which is essentially an expanding wave flowing through space-time. The team believes that these waves could help explain why some of the most distant galaxies out there appear to be more distant than they should be, according to the Standard Model of Cosmology (SM).

"We're saying that maybe these expanding waves are actually causing the anomalous acceleration. We're saying <u>dark energy</u> is not really the correct explanation," University of California in Davis (UCD) expert Blake Temple explains. The new set of equations revolves around Einstein's general theory of relativity, but also seems to offer a decent explanation for the observed cosmic expansion. Temple worked on the new calculations with University of Michigan colleague Joel Smoller, and the team published its results in the August 17th issue of the journal Proceedings of the National Academy of Sciences (PNAS).

"The research could change the way astronomers view the composition of our universe," the authors write in the summary of their journal entry, admitting, however, that more verifications are in order before a final conclusion is drawn. They also say that the new equations may prove to be a very potent alternative to dark energy theories simply because the latter were developed hastily, when astronomers discovered that the Universe was expanding at an ever-increasing speed, and had no explanation for this.

Dark energy "just seems like an unnatural correction to the equations – it's like a fudge factor. The equations don't make quite as much physical sense when you put it in. You just put it in to fit the data," Temple says, quoted by Space. "At this stage we think [the new equations are] a very plausible theory. We're saying there isn't any acceleration. The galaxies are displaced from where they're supposed to be because we're in the aftermath

ABSTRACT: In 1927, the American astronomer Edwin Hubble showed the Universe is expanding: distant galaxies are receding from each other. This confirmed the so-called *Standard Model of Cosmology*, that the universe, on the largest scale, is evolving according to a Friedman-Robertson-Walker spacetime. The starting assumption in this model is the *Cosmological Principle* that on the largest scale, we are not in a special place in the universe—that the universe is homogeneous and isotropic about every point like the FRW spacetime. In 1998, more accurate measurements of the recessional velocity of distant galaxies based on new Type 1a supernova data, made the astounding discovery that the Universe was actually *accelerating* relative to the standard model. So the Standard Model is incorrect. The explanation for the *Anomalous Acceleration of the Galaxies* is one of the great open problems of physics.

The only way to account for the Anomalous Acceleration and preserve the FRW framework and the Cosmological Principle is to modify the Einstein equations by adding a *Fudge Factor* called the *Cosmological Constant*. *Dark Energy*, the physical interpretation of the Cosmological Constant, is then an unknown source of anti-gravitation that, for the model to be correct, must account for some 70 percent of the energy density of the universe. This is stated as a fact on the NASA webpage. In this talk I introduce a new family of expanding wave perturbations of the Standard Model, and explore the possibility that these might account for the Anomalous Acceleration of the galaxies without the Cosmological Constant or Dark Energy. [Joint work with Joel Smoller] We prove that all of the self-similar spacetimes in the family are distinct from the non-critical  $k \neq 0$  Friedmann spacetimes thereby *characterizing* the critical k = 0 Friedmann universe as the unique spacetime lying at the intersection of these two one-parameter families.





• We show that the standard Friedmann Universe  $(k = 0, p = \frac{c^2}{3}\rho)$ can be extended to a 3-parameter family of exact non-interacting expansion waves in GR

 Removing a scaling law and imposing regularity at the center this reduces to a 1-parameter family of distict spacetimes that include the standard model, and introduce a correction to the Hubble constant

- Since non-interacting self-similar expansion waves represent possible timeasymptotic solutions in the theory of conservation laws:
- Q: Could corrections account for the anomalous acceleration of the galaxies w/o cosmological constant/dark energy?
- Q: A new set of solutions to test against the observations?

# INTRODUCTION TO COSMOLOGY





## Standard Model of Cosmology

• **1922** *Alexander Friedmann*:

Derived FRW solutions of the Einstein equations: 3-space of constant curvature expanding in time:

$$ds^{2} = -dt^{2} + R(t)^{2} \left\{ \frac{dr^{2}}{1 - kr^{2}} + r^{2} d\Omega^{2} \right\}$$

 The Big Bang theory based on the FRW metric was worked out by <u>George Lemaître</u> in the late 1920's leading to Hubble's comfirmation of redshift vs luminoscity consistent with an FRW spacetime

Hubble's Constant  $\equiv H \equiv \frac{\dot{R}}{R}$ 





The FRW metric when k=0:

• 
$$ds^2 = -dt^2 + R(t)^2 \left\{ dr^2 + r^2 d\Omega^2 \right\}$$

The universe is infinite flat space  $\mathbb{R}^3$  at each fixed time:

"Galaxies move along r = const., and  $\bar{r} = R(t)r$  measures distance at each fixed time"

 $\bigcirc$ 



• 
$$ds^2 = -dt^2 + R(t)^2 \left\{ dr^2 + r^2 d\Omega^2 \right\}$$

The universe is infinite flat space  $\mathbb{R}^3$  at each fixed time:

"E.g., in Standard Model, during radiation phase, after inflation..."

 $R(t) = \sqrt{t}$ 







Recent supernova data have tested the dependence of the Hubble constant on time, and the results don't fit standard model...

"Anomalous Acceleration of Galaxies"

Introduction of "Cosmological Const" and "Dark Energy"

Dark energy is non-classical Negative pressure Anti-gravity effect




The FRW Mathematical Model:

• Einstein Equations (1915): 
$$G_{ij} = \kappa T_{ij}$$

 $G_{ij}$ =Einstein Curvature Tensor

 $T_{ij} = (\rho + p)u_iu_j + pg_{ij}$ =Stress Energy Tensor (perfect fluid)

• Einstein Equations for k=0 Friedmann metric:

$$H^2 = \frac{\kappa}{3}\rho$$

$$\dot{\rho} = -3(\rho + p)H$$

**\*** Solutions determined by equation of state:  $p = p(\rho)$ 



#### Standard Model for Dark Energy







# The Question we Explore:

"Could the anomalous acceleration of the galaxies be due to the fact that we are looking outward into an expansion wave different from the k=0 FRW spacetime, and NOT due to a cosmological constant?"

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The Einstein equations have been confirmed without the cosmological constant in every setting except cosmology...

# The Question we Explore:

"Could the anomalous acceleration of the galaxies be due to the fact that we are looking out into an expansion wave different from the k=0 FRW spacetime, and NOT due to a cosmological constant?"

The Einstein equations have been confirmed without the cosmological constant in every setting except cosmology...

Note: A general expansion wave has a center of expansion...

The Einstein equations that describe the expansion of the Universe during the radiation phase of the expansion form a highly nonlinear system of coupled wave equations in the form of conservation laws. Such wave equations support the propagation of waves, and self-similar expansion waves are important because even when dissipative terms are neglected in conservation laws, the nonlinearities alone provide a mechanism whereby non-interacting selfsimilar wave patterns can emerge from general interactive solutions, via the process of wave interaction and shock wave dissipation



















To start: we proposed to numerically simulate the secondary reflected wave reflected back in our shock wave cosmology model...

#### **References:**

- Talk: Numerical Cosmology Session, National meeting, New Orleans, January 2007 http://www.math.ucdavis.edu/~temple/
- Thesis: numerical simulation by a locally inertial Godunov method, Zeke Volger, UC-Davis, 2009







 The numerical method required getting an explicit form for the (k = 0, p = 1/3 p)-FRW metric in Standard Schwarzschild Coordinates
 Upon doing this we found that there exists an integrating factor such that the metric satisfies an ODE in Standard Schwarzschild coordinates...the ODE's then introduce 3 extra free parameters...

...(the 3-initial conditions)!

Thus: we look for an expanding wave perturbation of the k=0 FRW metric during the period when:  $p = \frac{c^2}{3}\rho$  $ds^2 = -dt^2 + R(t)^2 \left\{ dr^2 + r^2 d\Omega^2 \right\}$  $\bigvee$  $R(t) = \sqrt{t}$  $H(t) = \frac{\dot{R}(t)}{R(t)} = \frac{1}{2t}$ 



A <u>Three</u> Parameter Family of Expanding Wave Solutions of the Einstein Equations including The Standard Model of Cosmology



 Spherically symmetric spacetime metrics can "generically" be mapped over to
 Standard Schwarzschild Coordinates... [c.f.Wein]

 In general there exist MANY ways to do this, depending on an INTEGRATING FACTOR that solves a PDE

$$\frac{\text{Theorem:}}{\text{metric}} \text{ Assume } p = \frac{c^2}{3}\rho, \ k = 0. \text{ Then the FRW}$$

$$\frac{ds^2 = -dt^2 + R(t)^2 dr^2 + \bar{r}^2 d\Omega^2,$$
under the mapping
$$\bar{r} = R(t)r,$$

$$\bar{t} = \left\{1 + \left[\frac{R(t)r}{2t}\right]^2\right\}t,$$
goes over to the SSC-metric
$$ds^2 = -\frac{d\bar{t}^2}{1 - v(\xi)^2} + \frac{d\bar{r}^2}{1 - v(\xi)^2} + \bar{r}^2 d\Omega^2,$$
where
$$\xi = \frac{\bar{r}}{\bar{t}} = \frac{2v}{1 + v^2}$$

**Corollary:** There exists a coordinate mapping that takes the  $p = \frac{1}{3}\rho$ , k = 0 FRW metric over to SSC-coordinates such that SSC metric components DEPEND ONLY ON THE SINGLE VARIABLE  $\xi = \frac{\overline{r}}{\overline{t}}$ (Like an expansion wave!) This implies that the standard FRW metric after inflation is equivalent to a metric that satisfies an ODE in SSC-Coordinates! We now construct this ODE systematically...








Q: When do	the SSC PDE's reduce to	ODE's?
Four	$\left\{-r\frac{A_r}{A} + \frac{1-A}{A}\right\} = \frac{\kappa B}{A}r^2T^{00}$	(1)
PDE's	$\frac{A_t}{A} = \frac{\kappa B}{A} r T^{01}$	(2)
	$\left\{ r\frac{B_r}{B} - \frac{1-A}{A} \right\} = \frac{\kappa}{A^2} r^2 T^{11}$	(3)
-	$\left\{ \left(\frac{1}{A}\right)_{tt} - B_{rr} + \Phi \right\} = 2\frac{\kappa B}{A}r^2T^{22},$	(4)
where		
$\Phi$ =	$\frac{B_t A_t}{2A^2 B} - \frac{1}{2A} \left(\frac{A_t}{A}\right)^2 - \frac{B_r}{r} - \frac{BA_r}{rA}$	
	$+\frac{D}{2}\left(\frac{D_r}{B}\right) -\frac{D}{2}\frac{D_r}{B}\frac{A_r}{A}.$	
Ans#I: A=A(r), B=B(r) time-independent		





...the fundamental equation of Newtonia astrophysics, with general-relativistic corrections supplied by the last three factors, [Weinberg, page 301].



## The SSC-equations reduce to ODE's when:

(I) 
$$T_{ij} = (\rho + p)u^i u^j + pg^{ij}$$
 is linear in  $\rho$ 

(2) 
$$A, B, v \text{ and } r^2 \rho \text{ depend on } \xi = r/t$$

Claim: one choice of initial conditions gives the standard model!



Four  
PDE's
$$\begin{cases}
-r\frac{A_r}{A} + \frac{1-A}{A} \\
= \frac{\kappa B}{A}r^2T^{00} \quad (1) \\
\frac{A_t}{A} = \frac{\kappa B}{A}rT^{01} \quad (2) \\
\{r\frac{B_r}{B} - \frac{1-A}{A} \\
= \frac{\kappa}{A^2}r^2T^{11} \quad (3) \\
-\{\left(\frac{1}{A}\right)_{tt} - B_{rr} + \Phi\} = 2\frac{\kappa B}{A}r^2T^{22}, \quad (4) \\
where
\end{cases}$$
where
$$\Phi = \frac{B_tA_t}{2A^2B} - \frac{1}{2A}\left(\frac{A_t}{A}\right)^2 - \frac{B_r}{r} - \frac{BA_r}{rA} \\
+ \frac{B}{2}\left(\frac{B_r}{B}\right)^2 - \frac{B}{2}\frac{B_r}{B}\frac{A_r}{A}.$$

$$(1)+(2)+(3)+(4) \qquad (1)+(3)+div T=0$$

Theorem: (Te-Gr) The equations close in a "locally inertial" formulation of (1), (2) & Div T=0: 
$$\begin{split} \left\{ T_M^{00} \right\}_{,0} + \left\{ \sqrt{AB} T_M^{01} \right\}_{,1} &= -\frac{2}{r} \sqrt{AB} T_M^{01}, \\ \left\{ T_M^{01} \right\}_{,0} + \left\{ \sqrt{AB} T_M^{11} \right\}_{,1} &= -\frac{1}{2} \sqrt{AB} \left\{ \frac{4}{r} T_M^{11} + \frac{(1-A)}{Ar} (T_M^{00} - T_M^{11}) \right\}_{,1} \end{split}$$
(1)(2) $+\frac{2\kappa r}{A}(T_M^{00}T_M^{11}-(T_M^{01})^2)-4rT^{22}\bigg\},$  $rA_r = (1-A) - \kappa r^2 T_M^{00},$ (3) $rB_r = \frac{B(1-A)}{4} + \frac{B}{4}\kappa r^2 T_M^{11}.$ (4) $T_M^{00} = \frac{c^4 + \sigma^2 v^2}{c^2 - v^2} \rho$  $T^{00} = \frac{1}{B}T_M^{00}$  $p = \sigma \rho \qquad \qquad T_M^{01} = \frac{c^2 + \sigma^2}{c^2 - v^2} cv \rho$  $T^{01} = \sqrt{\frac{A}{B}}T_M^{01}$  $T^{11} = AT_M^{11}$  $T_M^{11} = \frac{v^2 + \sigma^2}{c^2 - v^2} \rho c^2$  $v = \frac{1}{\sqrt{AB}} \frac{u^1}{u^0}$ 





•Substituting  $S^{ij}$ , (1), (2) & (3) become:  $rA_r = (1-A) - S^{00}$  (1)  $rA_t = \sqrt{AB} S^{01}$  (2)  $rB_r = \frac{B}{A} \{(1-A) + S^{11}\}$  (3) • Now assume  $A, B, S^{ij}$  depend only on  $\xi = \frac{r}{t}$   $A = A(\xi), \quad B = B(\xi), \quad S^{ij} = S^{ij}(\xi)$ • Then (1), (2) & (3) all reduce to ODE's in  $\xi$ !

 (1), (2) & (3) reduce to ODE's in  $\xi$ ...

$$\xi A_{\xi} = (1-A) - \kappa S^{00}$$
 (I)

$$\xi^2 A_{\xi} = \sqrt{AB} \kappa S^{01}$$
 (2)

$$\xi B_{\xi} = \frac{B}{A} \{ (1-A) + \kappa S^{11} \}$$
 (3)

• Equations (1) & (2) require the compatibility condition

$$(1-A) - \kappa S^{00} = \frac{\sqrt{AB}}{\xi} \kappa S^{01}$$



• A similar reduction applies to Equation (4):  

$$\{T_{M}^{01}\}_{,0} + \{\sqrt{AB}T_{M}^{11}\}_{,1} = -\frac{1}{2}\sqrt{AB}\left\{\frac{4}{r}T_{M}^{11} + \frac{(1-A)}{Ar}(T_{M}^{00} - T_{M}^{11}) + \frac{2\kappa r}{A}(T_{M}^{00}T_{M}^{11} - (T_{M}^{01})^{2}) - 4rT^{22}\right\}$$

• A similar reduction applies to Equation (4):  $\{T_{M}^{01}\}_{,0} + \{\sqrt{AB}T_{M}^{11}\}_{,1} = -\frac{1}{2}\sqrt{AB}\left\{\frac{4}{r}T_{M}^{11} + \frac{(1-A)}{Ar}(T_{M}^{00} - T_{M}^{11}) \quad (4) + \frac{2\kappa r}{A}(T_{M}^{00}T_{M}^{11} - (T_{M}^{01})^{2}) - 4rT^{22}\right\}$ • Multiplying through by  $r^{3}$  and using (\*) to eliminate w and  $w_{\xi}$  in favor of v we obtain (After considerable computation!)

• A similar reduction applies to Equation (4):  

$$\{T_{M}^{01}\}_{,0} + \{\sqrt{AB}T_{M}^{11}\}_{,1} = -\frac{1}{2}\sqrt{AB}\left\{\frac{4}{r}T_{M}^{11} + \frac{(1-A)}{Ar}(T_{M}^{00} - T_{M}^{11}) \quad (4) + \frac{2\kappa r}{A}(T_{M}^{00}T_{M}^{11} - (T_{M}^{01})^{2}) - 4rT^{22}\right\}$$
• Multiplying through by  $r^{3}$  and using (\*) to eliminate  $w$  and  $w_{\xi}$  in favor of  $v$  we obtain
$$(4) \quad \bigstar \quad \{v_{\xi} = -\left(\frac{1-v^{2}}{2\{\cdot\}_{D}}\right)\left\{(3+v^{2})G - 4v + \frac{4\left(\frac{1-A}{A}\right)\{\cdot\}_{N}}{(3+v^{2})G - 4v}\right\}$$

$$\{\cdot\}_{N} = \{-2v^{2} + 2(3-v^{2})vG - (3-v^{4})G^{2}\}$$

$$\{\cdot\}_{D} = \{(3v^{2} - 1) - 4vG + (3-v^{2})G^{2}\}$$

• A similar reduction applies to Equation (4):  

$$\{T_{M}^{01}\}_{,0} + \{\sqrt{AB}T_{M}^{11}\}_{,1} = -\frac{1}{2}\sqrt{AB}\left\{\frac{4}{r}T_{M}^{11} + \frac{(1-A)}{Ar}(T_{M}^{00} - T_{M}^{11}) \quad (4) + \frac{2\kappa r}{A}(T_{M}^{00}T_{M}^{11} - (T_{M}^{01})^{2}) - 4rT^{22}\right\}$$
• Multiplying through by  $r^{3}$  and using (\*) to eliminate  $w$  and  $w_{\xi}$  in favor of  $v$  we obtain
$$(4) \iff \qquad \xi v_{\xi} = -\left(\frac{1-v^{2}}{2\{\cdot\}_{D}}\right)\left\{(3+v^{2})G - 4v + \frac{4\left(\frac{1-A}{A}\right)\{\cdot\}_{N}}{(3+v^{2})G - 4v}\right\}$$

$$\begin{cases} \xi_{VAB} = \frac{\xi}{\sqrt{AB}} \\ \{\cdot\}_{D} = \{-2v^{2} + 2(3-v^{2})vG - (3-v^{4})G^{2}\} \\ \{\cdot\}_{D} = \{(3v^{2} - 1) - 4vG + (3-v^{2})G^{2}\} \end{cases}$$

Conclude: (1) = (2), (3), & Div T^{j1} = 0  
are Equivalent to:  

$$\xi A_{\xi} = -\left[\frac{4(1-A)v}{(3+v^2)G-4v}\right]$$
(1)  
(ODE) 
$$\xi G_{\xi} = -G\left\{\left(\frac{1-A}{A}\right)\frac{2(1+v^2)G-4v}{(3+v^2)G-4v}-1\right\}$$
(2)  

$$\xi v_{\xi} = -\left(\frac{1-v^2}{2\{\cdot\}_D}\right)\left\{(3+v^2)G-4v+\frac{4\left(\frac{1-A}{A}\right)\{\cdot\}_N}{(3+v^2)G-4v}\right\}$$
(3)  

$$\{\cdot\}_N = \left\{-2v^2+2(3-v^2)vG-(3-v^4)G^2\right\}$$

$$\{\cdot\}_D = \left\{(3v^2-1)-4vG+(3-v^2)G^2\right\}$$

$$\boxed{G = \frac{\xi}{\sqrt{AB}} ; \xi = \frac{r}{t}}$$

$$\begin{cases} \xi A_{\xi} = -\left[\frac{4(1-A)v}{(3+v^2)G-4v}\right] & (\mathbf{I}) \\ \xi G_{\xi} = -G\left\{\left(\frac{1-A}{A}\right)\frac{2(1+v^2)G-4v}{(3+v^2)G-4v}-1\right\} & (\mathbf{2}) \\ \xi v_{\xi} = -\left(\frac{1-v^2}{2\left\{\cdot\right\}_D}\right)\left\{(3+v^2)G-4v+\frac{4\left(\frac{1-A}{A}\right)\left\{\cdot\right\}_N}{(3+v^2)G-4v}\right\} & (\mathbf{3}) \end{cases}$$

A system of 3 ODE's analagous to the Oppenheimer-Volkoff Equations except they describe GR-Expansion Waves! **Theorem:** Assume that  $A(\xi)$ ,  $G(\xi)$  and  $v(\xi)$  solve ODE and use the constraint  $\kappa w \equiv \frac{r^2 \rho}{3(1-v^2)} = \frac{(1-A)G}{(3+v^2)G-4v}$ to define  $\rho$   $\rho = \frac{1}{\kappa} \frac{3(1-v^2)(1-A)G}{(3+v^2)G-4v} \frac{1}{\bar{r}^2}$ Then the metric  $ds^2 = -B(\xi)d\bar{t}^2 + \frac{1}{A(\xi)}d\bar{r}^2 + \bar{r}^2d\Omega^2$ solves the Einstein equations with equation of state  $p = \rho c^2/3.$  • The Result: a system of three ODE's plus one constraint equivalent to the Einstein equations assuming A, B, v and  $r^2\rho$  depend only on  $\xi = \frac{r}{t}$ :

$$\xi \begin{pmatrix} A \\ E \\ v \end{pmatrix}_{\xi} = F \begin{pmatrix} A \\ E \\ v \end{pmatrix} \qquad (2)$$

$$(3)$$

$$(4)$$

$$\kappa w = \frac{1 - A}{3 + v^2 - 4vE} \qquad (1)=(2)$$

• The equations for a three parameter family of  

$$GR-expansion waves$$

$$\begin{cases} \xi A_{\xi} = -\left[\frac{4(1-A)v}{(3+v^2)G-4v}\right] \quad (1) \\ \xi G_{\xi} = -G\left\{\left(\frac{1-A}{A}\right)\frac{2(1+v^2)G-4v}{(3+v^2)G-4v}-1\right\} \quad (2) \\ \xi v_{\xi} = -\left(\frac{1-v^2}{2\left\{\cdot\right\}_D}\right)\left\{(3+v^2)G-4v+\frac{4\left(\frac{1-A}{A}\right)\left\{\cdot\right\}_N}{(3+v^2)G-4v}\right\} \quad (3) \\ \left\{\cdot\right\}_N = \left\{-2v^2+2(3-v^2)vG-(3-v^4)G^2\right\} \\ \left\{\cdot\right\}_D = \left\{(3v^2-1)-4vG+(3-v^2)G^2\right\} \\ \kappa w = \frac{(1-A)G}{(3+v^2)G-4v} \quad (\text{Compatibility Constraint}) \quad (4) \end{cases}$$





Proof: Coordinate mapping IMPLIES:  

$$A = 1 - v^{2}, E = \frac{1}{\psi_{0}\xi}, \xi = \frac{2v}{\psi_{0}(1 + v^{2})}, v_{\xi} = \frac{\psi_{0}(1 + v^{2})^{2}}{2(1 - v^{2})}$$
Plug in and check:  

$$\begin{cases} \xi A_{\xi} = -\left[\frac{4(1 - A)v}{(3 + v^{2})G - 4v}\right] \qquad (1) \\ \xi G_{\xi} = -G\left\{\left(\frac{1 - A}{A}\right)\frac{2(1 + v^{2})G - 4v}{(3 + v^{2})G - 4v} - 1\right\} \qquad (2) \\ \xi v_{\xi} = -G\left\{\left(\frac{1 - v^{2}}{2\{\cdot\}_{D}}\right)\left\{(3 + v^{2})G - 4v + \frac{4(\frac{1 - A}{A})\{\cdot\}_{N}}{(3 + v^{2})G - 4v}\right\} \qquad (3) \\ \{\cdot\}_{N} = \{-2v^{2} + 2(3 - v^{2})vG - (3 - v^{4})G^{2}\} \\ \{\cdot\}_{D} = \{(3v^{2} - 1) - 4vG + (3 - v^{2})G^{2}\} \\ \kappa w = \frac{(1 - A)G}{(3 + v^{2})G - 4v} \qquad (Compatibility Constraint) \qquad (4) \end{cases}$$
"A surprisingly long calculation!"



$$(\bigcirc = \left(-4v + \frac{1+v^2}{2v}(1+3v^2)\right)\frac{2(1+v^2)}{(1-v^2)^2}$$

$$+$$

$$(\bigcirc = 2\left(-2+3\frac{1+v^2}{2}\right)\frac{v(1+v^2)}{(1-v^2)}$$

$$+$$

$$(\bigcirc = -4\frac{2(1+v^2)}{(1-v^2)^2}v^3$$

$$+$$

$$(\oiint = 4\frac{(1+v^2)^3}{2v(1-v^2)^2}$$

$$+$$

$$(\circlearrowright = -2\frac{1+v^2}{2v}(1-v^2)$$
The sum is equal to zero!

Conclude: The standard model of cosmology after inflation represents one solution of our ODE's corresponding to one initial condition...



Since the standard model represents 1-point in a 2-parameter family, we look for leading order corrections to the standard model determined from the nearby GR-expansion waves Linearizing about the center  $\xi = 0$ :

• One eigen-family tends to infinity as  $\xi \to 0$ 

• Two eigen-solutions stay finite as 
$$\xi \to 0$$
 and:  
 $A \to 1, B \to 1, v \to 0$ 

(One parameter is the scaling law...)

 Conclude: There is a smooth I-parameter family of distinct spacetimes that extend the standard model!



The following Theorem shows:

"Nearby solutions stay surprising close to FRW..." **Theorem:** There exist positive constants  $(\psi_0, a)$  such that the following estimates hold near  $\xi = 0$ .

$$v(\xi) = v_1(\xi) + \frac{(1-a^2)}{8}\psi_0^3\xi^3 + O(1)|a-1|\xi^4$$
  

$$A(\xi) = 1 - \frac{a^2\psi_0^2}{4}\xi^2 + O(1)|a-1|\xi^4$$
  

$$G(\xi) = \psi_0\xi + O(1)|a-1|\xi^5$$
  

$$\sqrt{AB} = \frac{1}{\psi_0} + O(1)|a-1|\xi^4$$
**Theorem 1:** To leading order in  $\xi$ , the 1-parameter family that extends the standard model of cosmology is given in SSC's by

$$ds^{2} = -\frac{d\bar{t}^{2}}{\psi_{0}^{2}\left(1 - \frac{a^{2}\psi_{0}^{2}\xi^{2}}{4}\right)} + \frac{d\bar{r}^{2}}{\left(1 - \frac{a^{2}\psi_{0}^{2}\xi^{2}}{4}\right)} + \bar{r}^{2}\Omega^{2}$$
$$\underbrace{v = \frac{\psi_{0}}{2}\xi}{\xi = \frac{\bar{r}}{\bar{t}}}$$
$$(a = 1) \equiv \text{Standard Model}$$

**Theorem 1:** To leading order in  $\xi$ , the 1-parameter family that extends the standard model of cosmology is given in SSC's by







Since the velocity field is  $\approx$ independent of "a", it follows that the <u>inverse</u> mapping from Standard Model to SSC's provides **a co-moving coordinate system** to leading order in  $\xi$  $\bar{r}(t,r) = \sqrt{t} r$  $\bar{t}(t,r) = \psi_0 \left(1 + \frac{r^2}{4}\right) t$ 









Conclude: an observer at the center would measure a fractional correction to the Hubble constant on the order of...

$$\Delta_a \equiv \frac{H_a - H}{H} = \frac{3}{8}(1 - a^2)\zeta^2 + O\left(|a^2 - 1|\zeta^4\right)$$

$$\zeta \equiv \frac{\bar{r}}{ct} \approx \frac{\bar{r}}{(c/H)} \approx \frac{Dist}{Hubble \ Length}$$

 $\approx$  "Fractional Distance from Center to Furthest Visible Objects" Moreover: using co-moving coordinates, we can calculate the leading order <u>correction</u> to the redshift vs luminosity relation as measured by an observer at the center of the spacetime:







LET:  

$$d_{\ell} \equiv Luminosity \ Distance = \left(\frac{L}{4\pi\ell}\right)^{1/2}$$
  
 $L = Absolute \ Luminosity = \frac{\text{Energy Emitted by Source}}{\text{Time}}$   
 $\ell \equiv Apparent \ Luminosity = \frac{\text{Power Recieved}}{\text{Area}}$   
 $z = \frac{\lambda_0}{\lambda_e} - 1 = Redshift \ Factor$ 



The redshift vs luminosity relation as measured by an observer at the center of the spacetime is given by:

$$d_{\ell} = 2t_0 z \left\{ 1 + \frac{a^2 - 1}{2} z + \cdots \right\} + H.O.T$$

The redshift vs luminosity relation as measured by an observer at the center of the spacetime is given by:

$$d_{\ell} = 2t_0 z \left\{ 1 + \frac{a^2 - 1}{2} z + H.O.T \right\}$$

...Quadratic correction quoted in PNAS...

The redshift vs luminosity relation as measured by an observer at the center of the spacetime is given by:  $d_{\ell} = 2t_0 z \left\{ 1 + \frac{a^2 - 1}{2} z + \frac{(a^2 - 1)(6a^2 + 13)}{6} z^2 \right\}$ +H.O.T



The redshift vs luminosity relation as measured by an observer at the center of the spacetime is given by:

$$d_{\ell} = 2t_0 z \left\{ 1 + \frac{a^2 - 1}{2} z + \frac{(a^2 - 1)(6a^2 + 13)}{6} z^2 \right\}$$
  
+H.O.T

(The calculation is nontrivial, and relies on simplifying features of the spacetime metric of the nearby expanding wave solutions...)







After the radiation phase:

The redshift vs luminosity relation evolves continuously with time

Therefore...

We conclude (by continuity) corrections to the redshift vs luminosity relation observed after the radiation phase of the Big Bang can be accounted for, at the leading order quadratic level, by adjustment of the free parameter "a". The next order correction is a VERIFIABLE PREDICTION of the model!!

(Work in progress)

A different coord. mapping casts new metric in a different light:

A different coord. mapping casts new metric in a different light:

$$\bar{r}(t,r) = \frac{t^a}{2} r$$
$$\bar{t}(t,r) = \psi_0 \left(1 + \frac{a^2 \zeta^2}{4}\right) t$$

A different coord. mapping casts new metric in a different light:

$$\begin{aligned} \bar{r}(t,r) &= t^{a/2} r\\ \bar{t}(t,r) &= \psi_0 \left(1 + \frac{a^2 \zeta^2}{4}\right) t \end{aligned}$$

$$\begin{aligned} & \checkmark \end{aligned}$$

$$ds^2 = -dt^2 + t^a dr^2 + \bar{r}^2 d\Omega^2 + a(1-a)\zeta dt d\bar{r} \end{aligned}$$



"In Fact: In these coordinates... metric is exactly flat 3-space at each fixed t=const ...just like the standard model..."  $ds^{2} = -dt^{2} + t^{a} \{ dr^{2} + r^{2}d\Omega^{2} \} + a(1-a)\zeta dt d\bar{r}$  A "Conservation Law" Scenario of the Big Bang w/o Cosmological Constant:

- Conservation Laws Decay to Non-interacting Time-Asymptotic Wave Patterns.
- After inflation, Universe is nearly flat, but due to errors, it decays by the nonlinearities of the radiation phase  $a \neq 1$ to a nearby non-interacting expansion wave
- We happen to be near the center of expansion, so looking outward, we observe a critical FRW with a small correction

The Lesson of Conservation Laws...

"Expansion waves and shock waves are fundamental to conservation laws, because even when dissipative terms are neglected, shock-wave dissipation by itself causes non-interacting wave patterns to emerge from interactive solutions" "I.e. The one fact most certain about the Standard Model is an early hot dense epoch in which all energy was radiation..." "...one might reasonably conjecture that decay to a non-interacting expanding wave might have occured (locally??) during the radiation phase due to the large nonlinearities associated with the large sound speed when  $p = \frac{c^2}{3}\rho$ ."


To make a testable prediction, we need to get the corrections at t=379,000 yrs, propagate errors with p=0 to present time, and look for the best fit.



Note: The expansion wave may not propagate as self-similar AFTER the radiation phase! We Like:

- This correction to the Hubble Constant is not put in "Ad Hoc"...
- It is derived from first principles starting from a theory of

**Expansion Waves** 

We Wonder:

- What scale might such expanding waves exist on...?
- Is there an inconsistency with WMAP Data...?
- Can this be accounted for in some inflationary scenario...?

Final Comment: These expanding waves near k=0 FRW represent a sort of "instability" in the Standard Model...

Thus: Even if they do not account for the anomalous acceleration...

One Has to Wonder why the Universe would choose a=1, k=0, FRW, and not one of these nearby non-interacting

**Expansion Waves?** 



http://www.foxnews.com/story/0,2933,430943,00.html

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/ Exist in Giant Cosmic Bubble – Science News   Science & Technology   Technology News 10/1/08 1:08 PM			
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RELATED	Until now, there has been no good way to choose between dark energy or the void explanation, but a new study outlines a potential test of the bubble scenario.		
Mysterious New 'Dark Flow' Discovered in Space Huge Planetary Collision Left Tons of Space Debris Solar Wind at Lowest Levels in 50 Years Our Sun May Be a 'Galactic Hitchhiker' Nearly Invisible Galaxy Found Orbiting Milky Way	If we were in an unusually sparse area of the universe, then things could look farther away than they really are and there would be no need to rely on dark energy as an explanation for certain astronomical observations.		
	"If we lived in a very large under-density, then the space- time itself wouldn't be accelerating," said researcher Timoth Clifton of Oxford University in England. "It would just be that the observations, if interpreted in the usual way, would look like they were."		
ADVERTISEMENTS Mortgage Rates Drop - Fec at 2% \$200,000 loan for \$708/month Free Quotes - No SSN Rqd. Save \$1000s! Mortgage.RefinanceFrontier.co	Scientists first detected the acceleration by noting that distant supernovae seemed to be moving away from us faster than they should be.		
	One type of supernova (called Type Ia) is a useful distance indicator, because the explosions always have the same intrinsic brightness.		
Buy a link her	Since light gets dimmer the farther it travels, that means that when the supernovae appear faint to us, they are far away, and when they appear bright, they are closer in.		

Until now, there has been no good way to choose between dark energy or the void explanation, but a new study outlines a potential test of the bubble scenario.

If we were in an unusually sparse area of the universe, then things could look farther away than they really are and there would be no need to rely on dark energy as an explanation for certain astronomical observations.

"If we lived in a very large under-density, then the spacetime itself wouldn't be accelerating," said researcher Timothy Clifton of Oxford University in England. "It would just be that the observations, if interpreted in the usual way, would look like they were."

Scientists first detected the acceleration by noting that distant supernovae seemed to be moving away from us faster than they should be.

## According to them...

Center	$\approx$	$15 \ MPC$
	$\approx$	50 Million Light Years
	$\approx$	Distance between
		clusters of galaxies
	$\approx$	$1/200 \ Distance \ Across$
		Visible Universe

## According to them...

Our view...

"Modeling an under-density during the p=0 stage can only model evolution after the wave has formed, but cannot give an explanation for the creation of such a wave..."

p=0 is "non-interacting"

# Conclude:

We are exploring the possibility that these expanding waves might provide a quantitative explanation for the formation of such an underdensity...

### Comparison of dark energy models: A perspective from the latest observational data

Miao Li,<sup>1,2,\*</sup> Xiao-Dong Li,<sup>3,2,†</sup> and Xin Zhang<sup>4,1,‡</sup>

<sup>1</sup>Kavli Institute for Theoretical Physics China, Chinese Academy of Sciences, Beijing 100190, China
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<sup>4</sup>Department of Physics, College of Sciences, Northeastern University, Shenyang 110004, China

In this paper, we compare some popular dark energy models with the assumption of a flat universe by using the latest observational data including the type Ia supernovae Constitution compiliation, the baryon acoustic oscillation measurement from the Sloan Digital Sky Survey and the Two Degree Field Galaxy Redshift Survey, and the cosmic microwave background measurement given by the five-year Wilkinson Microwave Anisotropy Probe observations. Model comparison statistics such as the Bayesian and Akaike information criteria are applied to assess the worth of the models. These statistics favor models that give a good fit with fewer parameters. Based on this analysis, we find that the simplest cosmological constant model that has only one free parameter is still preferred by the current data. For other dynamical dark energy models, we find that some of them, such as the *a* dark energy, constant *w*, generalized Chaplygin gas, and holographic dark energy models, can provide good fits to the current data, and three of them, numly, the agegraphic dark energy, Dvali-Gabadadze-Porrati, and Ricci dark energy models, are clearly disfavored by the data.

### I. INTRODUCTION

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Dark energy has become one of the most important issues of the modern cosmology ever since the observations of type Ia supernovae (SNe Ia) first indicated that the universe is undergoing an accelerated expansion at the present stage [1]. However, hitherto, we still know little about dark energy. The limited information we know about dark energy includes: it causes the cosmic acceleration; it accounts for two-thirds of the cosmic energy density; it is gravitationally repulsive; it does not appear to cluster in galaxies; and so on. Many cosmologists suspect that the identity of dark energy is the cosmological constant that fits the observational data well. While, one also has reason to dislike the cosmological constant since it always suffers from the theoretical problems such as the "fine-tuning" and "cosmic coincidence" puzzles [2]. The fine-tuning problem, also known as the "old cosmological constant problem," is motivated by the enormous discrepancy between the theoretical prediction for the cosmological constant and its measured value. The so-called "new cosmological constant problem," namely, the cosmic coincidence problem, questions why we just live in an era when the densities of dark energy and matter are almost equal, which also indicates that the cosmological constant scenario may be incomplete. Thus, a variety of proposals for dark energy have emerged.

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The possibility that dark energy is dynamical, for example, in a form of some light scalar field [3], has been explored by cosmologists for a long time. A basic way to explore such a dynamical dark energy model in light of observational data is to parameterize dark energy by an equation-of-state parameter w, relating the dark energy pressure p to its density  $\rho$  via  $h = w\rho$ . In general, this parameter w is time variable. The most commonly used forms of w(a) involve the constant equation of state, w = const., and the Chevalliear-Polarski-Linder form [4],  $w(a) = w_0 + (1 - a)w_a$ , where  $w_0$  and  $w_a$  parameterize the present-day value of w and the first derivative. There are also many other dynamical dark energy models which stem from different aspects of new physics. For example, the "holographic dark energy" models [5, 6, 7, 8, 9, 10, 11] arise from the holographic principle of quantum gravity theory, and the Chaplygin gas models [12, 13, 14] are motivated by brane world scenarios and may be able to unify dark matter and dark energy. In addition, there is also significant interest in modifications to general relativity, in the context of explaining the acceleration of the universe. The Dvali-Gabadadze-Porrati models [15, 16, 17] arise from a class of brane-related theories in which gravity leaks out into the bulk at large distances, leading to the accelerated expansion of the universe.

In the face of so many competing dark energy candidates, it is important to find an effective way to decide which one is right, or at least, which one is most favored by the observational data. Although the accumulation of the current observational data has opened a robust window for constraining the parameter space of dark energy models, the model filtration is still a difficult mission owing to the accuracy of current data as well as the complication caused by different parameter numbers of various dark energy models. In this paper, we make an effort to assess some popular dark energy models in light of the latest observational



cosmological model. Needless to remind that even if a cosmological model is in agreement with all observations, whatever their accuracy, it does not prove that it is the "correct" model of the Universe, in the sense that it is the correct cosmological extrapolation and solution of the local physical laws.

Dark energy confronts us with a compatibility problem since, in order to "save the phenomena" of the observations, we have to include new ingredients (constant, matter fields or interactions) beyond those of our established physical theories. <u>However the required value for the simplest dark energy</u> model, i.e. the cosmological constant, is more than 60 order of magnitude smaller to what is expected from theoretical grounds (§ 1.1.6). This tension between what is required by astronomy and what is expected from physics reminds us of the twenty centuries long debate between Aristotelians and Ptolemeans (Duhem, 1913), that was resolved not only by the Copernican model but more important by a better understanding of the physics since



### References

- Exact solution incorporating a shock-wave into the standard FRVV metric for cosmology...
- Smoller-Temple, Shock-Wave Cosmology Inside a Black Hole, PNAS Sept 2003.
- Smoller-Temple, *Cosmology, Black Holes, and Shock Waves Beyond the Hubble Length,* Meth. Appl. Anal., 2004.

## References:

- Connecting the shock wave cosmology model with Guth's theory of inflation...
- How inflationary spacetimes might evolve into spacetimes of finite total mass, with J. Smoller, Meth. and Appl. of Anal., Vol. 12, No. 4, pp. 451-464 (2005).
- How inflation is used to solve the flatness problem, with J. Smoller, Jour. of Hyp. Diff. Eqns. (JHDE) Vol. 3, no. 2, 375-386 (2006).

	Comparison of dark energy models: A perspective from the latest observational data
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### Comparison of dark energy models: A perspective from the latest observational data

Miao Li,<sup>1,2,\*</sup> Xiao-Dong Li,<sup>3,2,†</sup> and Xin Zhang<sup>4,1,‡</sup>

<sup>1</sup>Kavli Institute for Theoretical Physics China, Chinese Academy of Sciences, Beijing 100190, China Laboratory of Frontiers in Theoretical Physics, Institute of Theoretical Physics, Chinese Academy of Sciences, Beijing 100190, <sup>3</sup>Interdisciplinary Center for Theoretical Study, University of Science and Technology of China, Hefei 230026, China <sup>4</sup>Department of Physics, College of Sciences, Northeastern University, Shenyang 110004, China

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Temple of the University of California, Davis. "We're saying that dark energy may not really be the correct explanation."

The researchers derived a set of equations describing expanding waves that fit Einstein's theory of general relativity, and which could also account for the apparent acceleration. Temple outlines the new idea with Joel Smoller of the University of Michigan in the Aug. 17 issue of the journal Proceedings of the National Academy of Sciences.

While more research will be needed to see if the idea holds up, "the research could change the way astronomers view the composition of our universe," according to a summary from the journal.

To convince other cosmologists, the new model will have to pass muster with further inquiry.

"There are many observational tests of the standard cosmological model that the proposed model must pass, aside from the late phase of accelerated expansion," said Avi Loeb, director of the Institute for Theory and Computation at the Harvard-Smithsonian Center for Astrophysics. "These include big bang nucleosynthesis, the quantitative details of the microwave background anisotropies, the Lyman-alpha forest, and galaxy surveys. The authors do not discuss how their model compares to these tests, and whether the number of free parameters they require in order to fit these observational constraints is smaller than in the standard model. Until they do so, it is not clear why this alternative model should be regarded as advantageous."

Johns Hopkins University astrophysicist Mario Livio agreed that to be seriously considered, the model must be able to predict properties of the universe that astronomers can measure.

He said the real test "is in whether they are able to reproduce all the observed cosmological parameters (as determined, e.g. by a combination of the Hubble Constant and the parameters determined by the CMB observations). To only produce an apparent acceleration is in itself interesting, but not particularly meaningful."

http://www.space.com/scienceastronomy/090817-dark-energy-alternative.html#comments

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Big Brain Theory: Have Cosmologists Lost Theirs? - New York Times

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The New Hork Times

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January 15, 2008

### **Big Brain Theory: Have Cosmologists Lost Theirs?**

### By <u>DENNIS OVERBYE</u>

Correction Appended

It could be the weirdest and most embarrassing prediction in the history of cosmology, if not science.

If true, it would mean that you yourself reading this article are more likely to be some momentary fluctuation in a field of matter and energy out in space than a person with a real past born through billions of years of evolution in an orderly star-spangled cosmos. Your memories and the world you think you see around you are illusions.

This bizarre picture is the outcome of a recent series of calculations that take some of the bedrock theories and discoveries of modern cosmology to the limit. Nobody in the field believes that this is the way things really work, however. And so in the last couple of years there has been a growing stream of debate and dueling papers, replete with references to such esoteric subjects as reincarnation, multiple universes and even the death of spacetime, as cosmologists try to square the predictions of their cherished theories with their convictions that we and the universe are real. The basic problem is that across the eons of time, the standard theories suggest, the universe can recur over

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Alan Guth, a cosmologist at the <u>Massachusetts Institute of Technology</u> who agrees this overabundance is absurd, pointed out that some calculations result in an infinite number of free-floating brains for every normal brain, making it "infinitely unlikely for us to be normal brains." Welcome to what physicists call the Boltzmann brain problem, named after the 19th-century Austrian physicist Ludwig Boltzmann, who suggested the mechanism by which such fluctuations could happen in a gas or in the universe.



